

Near-infrared spectroscopy of a nitrogen-loud quasar SDSS J1707+6443

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ABSTRACT

We present near-infrared spectroscopy of the $z \simeq 3.2$ quasar SDSS J1707+6443, obtained with MOIRCS on the Subaru Telescope. This quasar is classified as a “nitrogen-loud” quasar because of the fairly strong N III and N IV semi-forbidden emission lines from the broad-line region (BLR) observed in its rest-frame UV spectrum. However, our rest-frame optical spectrum from MOIRCS shows strong [O III] emission from the narrow-line region (NLR) suggesting that, at variance with the BLR, NLR gas is not metal-rich. In order to reconcile these contradictory results, there may be two alternative possibilities; (1) the strong nitrogen lines from the BLR are simply due to a very high relative abundance of nitrogen rather than to a very high BLR metallicity, or (2) the BLR metallicity is not representative of the metallicity of the host galaxy, better traced by the NLR. In either case, the strong broad nitrogen lines in the UV spectrum are not indication of a chemically enriched host galaxy. We estimated the black hole mass and Eddington ratio of this quasar from the velocity width of both C IV and H β , that results in $\log(M_{\text{BH}}/M_{\odot}) = 9.50$ and $\log(L_{\text{bol}}/L_{\text{Edd}}) = -0.34$. The relatively high Eddington ratio is consistent with our earlier result that strong nitrogen emission from BLRs is associated with high Eddington ratios. Finally, we detected significant [Ne III] emission from the NLR, implying a quite high gas density of $n_e \sim 10^6 \text{ cm}^{-3}$ and suggesting a strong coupling between quasar activity and dense interstellar clouds in the host galaxy.

Key words. Galaxies: active - Galaxies: nuclei - quasars: emission lines - quasars: individual: SDSS J1707+6443

1. Introduction

The chemical composition of galaxies is a powerful tool to distinguish various evolutionary scenarios as it is the result of past star-formation activity and is also affected by gas inflows and outflows. The chemical properties of active galactic nuclei (AGN) are particularly interesting, because the huge luminosities of AGN enable us to accurately measure spectroscopic features of quasars even at high redshifts and thus to explore the chemical properties in the early universe. Another interesting aspect of assessing the chemical composition of AGNs is that it provides clues on the co-evolution between galaxies and their supermassive black holes (SMBHs), that has been inferred by the tight correlation between the mass of SMBHs (M_{BH}) and their host spheroidal observed in the local Universe (e.g., Marconi & Hunt 2003; Ferrarese & Merritt 2000; Gebhardt et al. 2000). Matsuoka et al. (2011) reported a tight relationship between the metallicity in broad-line regions (BLRs) and M_{BH} at $z \sim 2 - 3$, suggesting that the evolution of SMBHs is associated with the cumulative star formation in the host galaxies (see also Warner et al. 2004; Nagao et al. 2006b).

A widely adopted method to estimate the BLR metallicity (Z_{BLR}) exploits the flux of nitrogen emission lines such as

N V $\lambda 1240$. Nitrogen is a secondary element and thus its relative abundance is proportional to the metallicity, $[\text{N}/\text{O}] \propto [\text{O}/\text{H}]$ or, equivalently, $[\text{N}/\text{H}] \propto [\text{O}/\text{H}]^2$ (see, e.g., Hamann & Ferland 1992). Based on this method, it has been found that the BLR in most quasars show super-solar metallicities (e.g., Nagao et al. 2006b), even for quasars at $z \gtrsim 6$ (Jiang et al. 2007; Juarez et al. 2009; Mortlock et al. 2011). The BLR metallicity is up to $Z_{\text{BLR}} \sim 10Z_{\odot}$ in the most extreme cases although it seems hard to understand such high metallicities through ordinary galaxy chemical evolution models (e.g., Hamann & Ferland 1993). More interestingly, some quasars show extremely strong emission in N V and in other nitrogen lines (especially N IV $\lambda 1486$ and N III $\lambda 1750$ semi-forbidden lines). The estimated Z_{BLR} reaches $\sim 15 Z_{\odot}$, or higher, and these quasars are classified as “N-loud quasars” (e.g., Baldwin et al. 2003; Bentz & Osmer 2004; Bentz et al. 2004). Here the definition of a N-loud quasar is that of Jiang et al. (2008); i.e., type 1 quasars with strong nitrogen emission of $EW_{\text{rest}}(\text{N IV } \lambda 1486) > 3 \text{ \AA}$ or $EW_{\text{rest}}(\text{N III } \lambda 1750) > 3 \text{ \AA}$ (see Section 2). Note that the measurement of the N V $\lambda 1240$ flux is sometimes uncertain due to the heavy blending with the Ly α emission (see, e.g., Nagao et al. 2006b); however the N-loudness of these quasars is convincing since the well-isolated N IV $\lambda 1486$ and N III $\lambda 1750$ are used for defining the N-loud

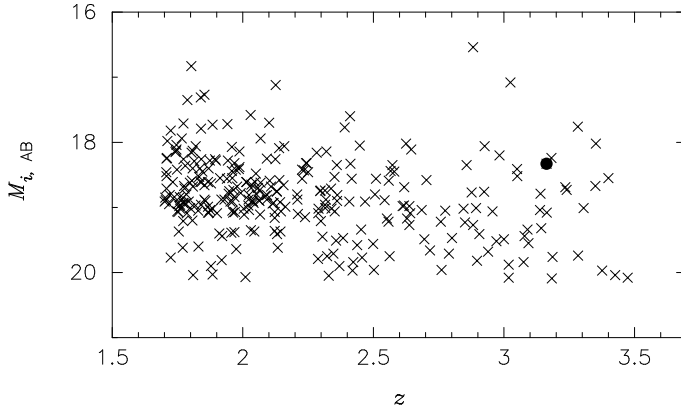


Fig. 1. Absolute i' -band magnitudes of 293 N-loud quasars, as a function of redshift. Filled circle and crosses denote SDSS J1707+6443 and the other N-loud quasars, respectively. Data are taken from Jiang et al. (2008).

quasar population. Standard chemical evolution models can not predict such extremely high metallicities at any epoch (e.g., Ballero et al. 2008), thus the presence of these N-loud quasars is a great challenge for galaxy evolutionary models.

Jiang et al. (2008) pointed out that N-loud quasars may not have such high Z_{BLR} , but simply have unusually high relative abundance of nitrogen in the BLR, mainly because the emission-line spectrum of N-loud quasars is not significantly different from that of typical quasars except for the nitrogen lines. This idea poses into question the use of the N v emission for Z_{BLR} measurements in quasars. It is therefore very important to verify observationally if N-loud quasars have either high Z_{BLR} or just high nitrogen relative abundances. More recently, Matsuoka et al. (2011) reported that the strength of nitrogen lines are correlated with the quasar Eddington ratio ($L_{\text{bol}}/L_{\text{Edd}}$) and discussed the possibility that quasars with strong nitrogen emission are in a special phase of the evolutionary history of their host galaxies. These studies outline the importance of understanding the nature of N-loud quasars, especially in terms of their chemical evolution. Therefore it is important to assess the metallicity of BLRs and host galaxies of N-loud quasars with a method different from that relying on rest-frame UV broad nitrogen lines.

In this paper, we focus on the emission of the narrow-line region (NLR) in N-loud quasars. Since strong nitrogen lines from their BLRs suggest a high Z_{BLR} , we examine the properties of the NLR to independently investigate the metallicity in the N-loud quasar population. While the BLR is located in very compact, subpc-scale regions around SMBHs, NLR gas clouds are distributed on kpc scales, i.e. on the scales of the host galaxies, thus enabling us to study the possible relationships between the properties of N-loud quasars and their host galaxies. Throughout this paper, we adopt a cosmology with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_m = 0.3$, and $\Omega_\Lambda = 0.7$.

2. Sample Selection, Observation, and the Data Reduction

In order to investigate NLR properties of N-loud quasars in detail, we focus on SDSS J1707+6443 at $z = 3.163$ from the N-loud quasar catalog of Jiang et al. (2008). This catalog contains 293 N-loud quasars selected from the Fifth Data Release quasar catalog (Schneider et al. 2007) of the Sloan Digital Sky Survey

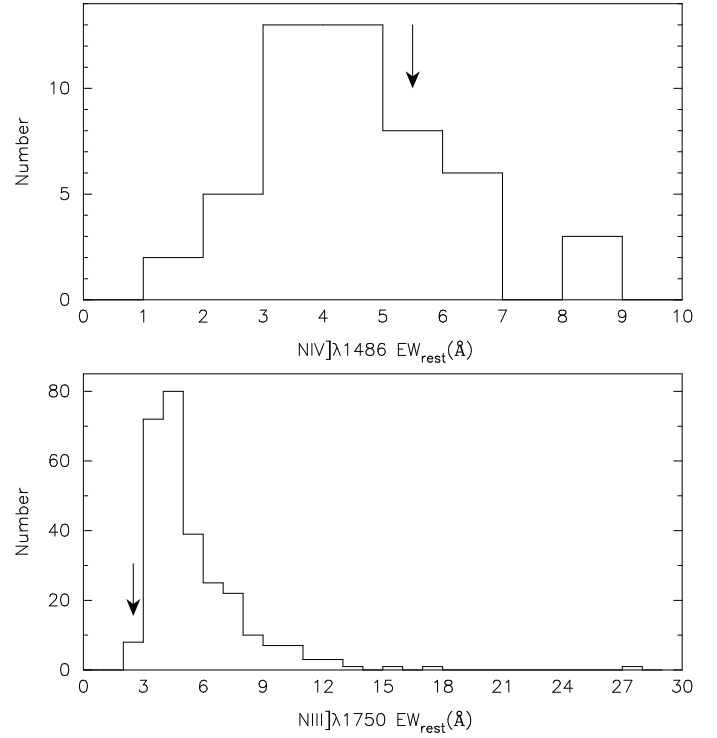


Fig. 2. The histograms of EW_{rest} of N iv] λ 1486 (upper panel) and N iii] λ 1750 (lower panel). The arrows denote the EW_{rest} values of SDSS J1707+6443. Data are taken from Jiang et al. (2008).

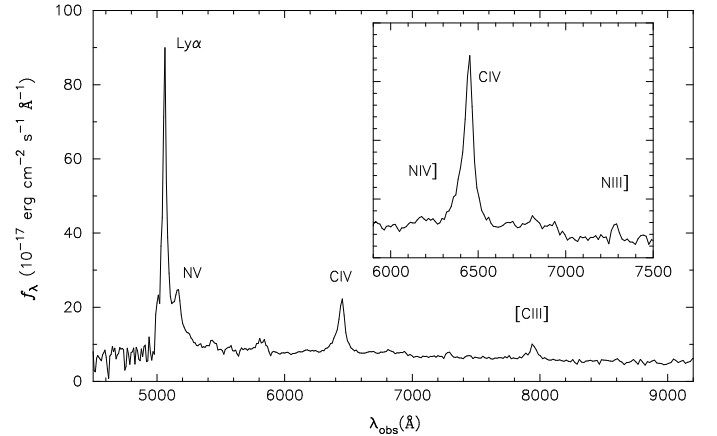


Fig. 3. Optical spectrum of SDSS J1707+6443, obtained from the SDSS database. Line IDs are given for strong emission lines. The inset shows spectral features around the C iv emission, where N iv] λ 1486 and N iii] λ 1750 lines are seen.

(SDSS; York et al. 2000). The N-loud quasars in the catalog of Jiang et al. (2008) were selected with the following criteria; (i) $i' < 20.1$, (ii) $1.7 < z < 4.0$, and (iii) $EW_{\text{rest}}(\text{N iv}]\lambda 1486) > 3 \text{ \AA}$ or $EW_{\text{rest}}(\text{N iii}]\lambda 1750) > 3 \text{ \AA}$. Among the N-loud quasars listed in the catalog of Jiang et al. (2008), we focused on SDSS J1707+6443 because it is relatively bright and its redshift is adequate for detecting redshifted NLR lines without suffering from the deep atmospheric absorption (Figure 1). In Figure 2, we show the frequency distributions of $EW_{\text{rest}}(\text{N iv}]\lambda 1486)$ and $EW_{\text{rest}}(\text{N iii}]\lambda 1750)$, and also show where SDSS J1707+6443 is located in these frequency distributions (see also Table 1). The averages and standard deviations of $EW_{\text{rest}}(\text{N iv}]\lambda 1486)$ distri-

Table 1. Observational properties of SDSS J1707+6443

name	z	i'_{AB} (mag)	$i'_{abs,AB}$ (mag) ^a	H_{Vega} (mag) ^b	EW_{rest} (Å) ^c			
					N IV]	N III]	C IV	C III]
SDSS J1707+6443	3.163	18.33	−28.06	16.15	5.1	2.5	40.8	15.3

^a i' band absolute magnitude given by Schneider et al. (2007).^b 2MASS magnitude given by Schneider et al. (2007).^c Rest-frame equivalent widths of some BLR lines given by Jiang et al. (2008).

butions are $4.6\text{Å} \pm 1.7\text{Å}$, with minimum and maximum values in the parent sample of 1.2Å and 8.9Å . The averages and standard deviations of $EW_{rest}(\text{N III}]\lambda 1750)$ are $5.6\text{Å} \pm 2.6\text{Å}$, with minimum and maximum values of 2.5Å and 27.1Å . Figure 2 does not suggest that SDSS J1707+6443 is located at the low-end in the distribution of $EW_{rest}(\text{N III}]\lambda 1750)$, because some of N-loud quasars do not always show detectable emission lines of both N IV] $\lambda 1486$ and N III] $\lambda 1750$. More specifically, only 50 N-loud quasars out of 293 show detectable N IV] $\lambda 1486$ emission while 280 out of 293 show detectable N III] $\lambda 1750$ emission (Jiang et al. 2008). Although SDSS J1707+6443 has a high N IV] $\lambda 1486$ /N III] $\lambda 1750$ ratio with respect to typical N-loud quasars, we do not discuss this issue in this paper since it is beyond the scope of this work. The SDSS spectrum and the basic observational properties of SDSS J1707+6443 are shown in Figure 3 and Table 1.

To assess the NLR properties of type 1 quasars, it is necessary to investigate rest-frame optical spectra and thus near-infrared spectroscopic data are required for this target. Therefore we observed SDSS J1707+6443 with MOIRCS (Ichikawa et al. 2006; Suzuki et al. 2008), the near-infrared spectrograph at the Subaru Telescope, on 30th May 2009. We used the HK grism, that results into the wavelength coverage $1.4\mu\text{m} \leq \lambda_{obs} \leq 2.3\mu\text{m}$. By using the $0.6''$ -wide slit, the resulting spectral resolution is $R \sim 570$, which is measured through the widths of the observed OH airglow emission. The total integration time was 1800 sec, consisting of six 300 sec independent exposures. We also observed HIP 86687 (A2 star with $H_{Vega} = 8.9$ and an assumed effective temperature of 8810 K) for the flux calibration and the telluric absorption correction. During the observation, the typical seeing size was $0.8''$ in the optical.

Standard data processing was performed using the available IRAF¹ routines. Frames were flat-fielded using domeflat images, and sky emission was performed by subtracting pairs of subsequent frames with the target at different positions along the slit. Cosmic-ray events were then removed by using the `lineclean` task. We extracted the one-dimensional spectrum of the target using the `apall` task, with an aperture size of ± 5 pixels ($1.17''$) from the emission center and, in this process, sky residuals were removed. Wavelength calibration was performed using OH sky lines. Finally, flux calibration and telluric absorption correction were carried out using the observed spectra of HIP 86687. The final processed spectrum was obtained by combining all single exposure frames.

To check the consistency of the flux calibration between the SDSS data and our MOIRCS data, we extrapolated the UV continuum emission toward the longer wavelength and compared the flux density at $\lambda_{rest} = 5100\text{Å}$ between the extrapolated spec-

trum and our MOIRCS spectrum. The adopted spectral index is derived by fitting the SDSS spectrum at the wavelength regions where strong emission-line features are not present ($\lambda_{rest} \sim 1350\text{Å}$, 1450Å , 1670Å , and 1970Å). Here we assume a constant spectral index at $\lambda_{rest} < 4000\text{Å}$ (see Vanden Berk et al. 2001). The discrepancy in the continuum flux is $\sim 10\%$, and therefore we conclude that the flux calibration is consistent between the SDSS spectrum and our MOIRCS spectrum.

3. Result

The final processed spectrum of SDSS J1707+6443 is shown in Figure 4, with the Mauna Kea atmospheric transmission² and a typical sky spectrum obtained during our MOIRCS observing run. A prominent broad H β emission is detected at $\lambda_{obs} \sim 2.0\mu\text{m}$, a typical spectral feature from quasar BLRs. In addition to the broad H β emission, some narrow emission lines from the NLR are also detected like the [O III] doublet in the K -band, a feature common in several high- z quasars (e.g., McIntosh et al. 1999; Netzer et al. 2004; Marziani et al. 2009; Greene et al. 2010). Interestingly, we also detect the quite rare [Ne III] $\lambda 3869$ emission line but not the [Ne V] $\lambda 3426$ and [O II] $\lambda 3727$ lines.

To measure the emission-line fluxes and velocity widths, we fit the spectral features by using the `specfit` routine (Kriss 1994). Here we adopt a single Gaussian profile for forbidden narrow lines and a double Gaussian profile for H β line. The measured emission-line properties are summarized in Table 2, where the presented quantities are based on the fitting models, not based on the actual data themselves. In Figure 5, we show fitting results and residuals. The measured velocity widths ($\Delta v_{FWHM} \sim 600\text{--}850\text{ km s}^{-1}$ for the NLR lines and $\Delta v_{FWHM} \sim 5500\text{ km s}^{-1}$ for the BLR line) are consistent with typical values for NLR and BLR emission lines seen in type 1 AGNs ($426 \pm 251\text{ km s}^{-1}$ and $4420 \pm 3210\text{ km s}^{-1}$ for the NLR and BLR lines; Ho & Kim 2009). Note that the systematic errors given for the emission-line widths are estimated by applying some fitting functions for emission lines (such as Gaussian, Lorentzian, and so on) and examining the standard deviation of the width. Here the statistical errors in the velocity width are smaller than the systematic errors, given the achieved signal-to-noise ratio.

4. Discussion

4.1. The Chemical Properties of the Narrow Line Region

One of the aims of this study is to estimate the metallicity of the host galaxy of a N-loud quasar, SDSS J1707+6443. More specifically, we wish to test whether this quasar has extremely metal-rich gas clouds, as expected by its strong broad nitrogen emission lines. Although strong broad nitrogen lines are

¹ IRAF is distributed by the National Optical Astronomy Observatory, which is operated by the Association of Universities for Research in Astronomy (AURA) under cooperative agreement with the National Science Foundation.

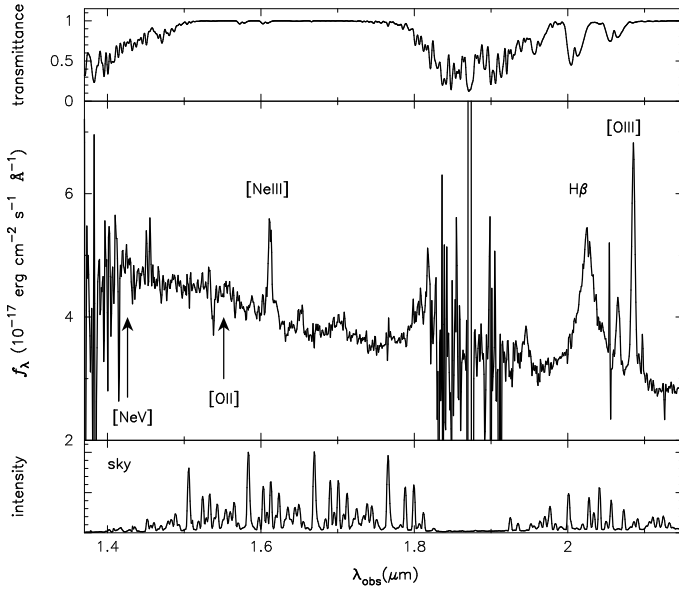
² Data obtained from the UKIRT web site.

Table 2. The measured spectral features

line ID	center wavelength [Å]	EW _{rest} [Å]	line flux [10 ⁻¹⁷ erg s ⁻¹ cm ⁻²]	FWHM ^a [Å]	FWHM ^b [km s ⁻¹]
[Ne v]λ3426	—	—	<12.3 ^c	—	—
[O II]λ3727	—	—	<10.0 ^c	—	—
[Ne III]λ3869	16117.5	4.9±0.2	83.3±3.7	54.5±6.5	866 ±103
Hβ _{narrow}	20245.7	4.1±0.2	51.8±2.3	55.3±8.4	627±95
Hβ _{broad}	20267.1	49.3±0.5	628.6±5.9	379.9±67.8	5595±999
[O III]λ4959	20653.8	5.8±0.1	72.5±1.7	56.4±8.6	627±96
[O III]λ5007	20853.7	17.7±0.1	217.9±1.7	57.0±8.7	627±96

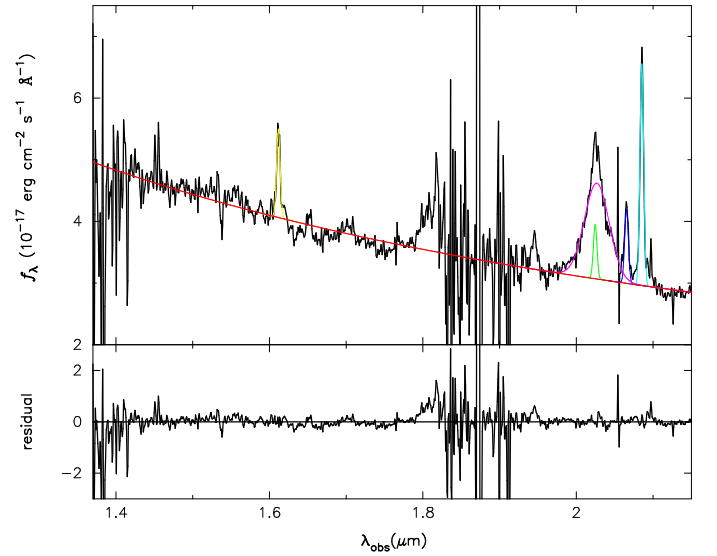
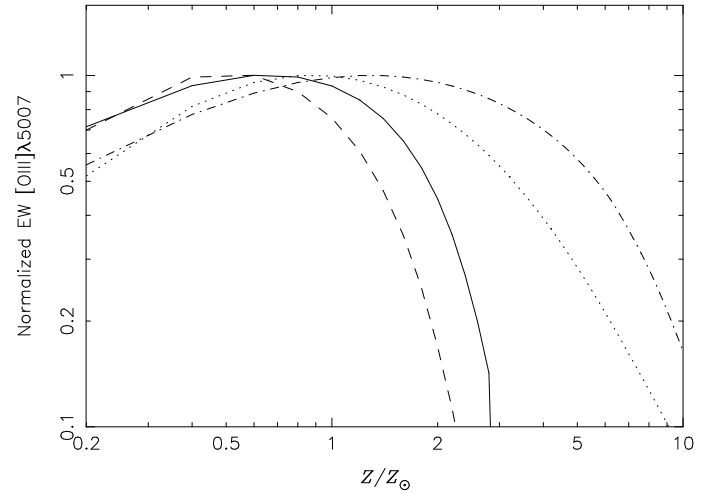
^a Measured emission-line width before the correction for the instrumental broadening.

^b Emission-line velocity width after the correction for the instrumental broadening.

^c For the undetected lines, 3σ upper-limit fluxes are given.

Fig. 4. Near-infrared spectrum of SDSS J1707+6443 obtained in our MOIRCS run (middle panel), with the Mauna Kea atmospheric transmission (upper panel) and the typical sky spectrum obtained during our MOIRCS run (lower panel). Important emission lines are labeled in the middle panel.

usually interpreted as indications of high Z_{BLR} (for both permitted lines and semi-forbidden lines; see, e.g., Shields 1976; Hamann & Ferland 1992), Jiang et al. (2008) pointed out that N-loud quasars might simply have unusually high nitrogen relative abundances. This is because the emission-line spectrum of the N-loud quasars is similar to that of usual type 1 quasars. In particular, many N-loud quasars do not show anomalous behavior for what concerns other broad emission-line flux ratios, not involving nitrogen, but sensitive to Z_{BLR} , such as $(\text{Si IV} + \text{O IV})/\text{C IV}$ (Nagao et al. 2006b; Juarez et al. 2009; Simon & Hamann 2010). However, the nuclear BLR involves only a very small fraction of the gas content in the galaxy and, therefore, may well not be representative of the metallicity in the host galaxy. As a consequence, it may be more instructive to investigate the metallicity in the NLR (Z_{NLR}) through the rest-frame optical spectrum, so to investigate whether the high metallicities inferred for the BLR are confirmed on the larger scales traced by the NLR.

We focus first on the metallicity dependence of line emissivity on the physical conditions typical of the NLR. The emissivity of collisionally excited emission lines strongly depends


Fig. 5. The fitting result (upper panel) and the residual (lower panel) for the processed MOIRCS spectrum. See the main text for details of the fitting procedure.

Fig. 6. Predicted equivalent width of [O III]λ5007, as a function of Z_{NLR} . Solid, dashed, dot-dashed, and dot lines denote the models with $(\log n_{\text{H}}, \log U) = (1, -1.5), (1, -3.5), (4, -1.5),$ and $(4, -3.5)$, respectively. The EW predictions are normalized by their peak values.

on the gas temperature. The equilibrium temperature of ionized gas clouds depends on gas metallicity, because metal emission lines are the main coolants of these clouds. Therefore, collisionally excited emission lines from the NLR become fainter for increasing Z_{NLR} (e.g., Nagao et al. 2006a; Matsuoka et al. 2009). To show this effect more explicitly, we computed simple photoionization model calculations by using Cloudy version 08.00 (Ferland et al. 1998; Ferland 2006). We assume ionization-bounded plane-parallel gas clouds with a constant density, photoionized by the typical spectral energy distribution for quasars (Mathews & Ferland 1987; Ferland 2006). We examine the metallicity dependence of the equivalent width (EW) of the [O III] emission, with the relative elemental abundance ratios fixed to the solar values except for nitrogen. We assume that the nitrogen relative abundance (N/H) scales with the square of the metallicity. In Figure 6, we show photoionization models results for gas densities of $n_{\text{H}} = 10^1 \text{ cm}^{-3}$ and 10^4 cm^{-3} , and ionization parameters of $U = 10^{-3.5}$ and $10^{-1.5}$. The emissivity curves in this figure are normalized by their peak to highlight their dependence on the gas metallicity. Clearly the $EW([\text{O III}])$ decreases rapidly as the metallicity increases, and the resultant $EW([\text{O III}])$ is quite small at $Z_{\text{NLR}} > 5 Z_{\odot}$. Note that we cannot derive accurate values of Z_{NLR} based only on [O III], since detailed metallicity studies requires multiple emission lines (e.g., Storch-Bergmann et al. 1998; Nagao et al. 2002; Iwamuro et al. 2003; Groves et al. 2006; Nagao et al. 2006a; Matsuoka et al. 2009). However, the aim of this paper is only to assess whether the NLR in SDSS J1707+6443 has a high metallicity, not to derive its accurate value. These results suggest that we can discriminate whether or not N-loud quasars have a relatively high Z_{NLR} with respect to other type 1 quasars by comparing the $EW([\text{O III}])$ distributions of the two populations.

In Figure 7, we show the frequency distribution of $EW_{\text{rest}}([\text{O III}])$ of type 1 quasars at $1 < z < 4$, that are compiled from the literature (McIntosh et al. 1999; Netzer et al. 2004; Marziani et al. 2009; Greene et al. 2010; shown with a solid histogram in Figure 7). In these four papers, there are 86 [O III]-detected quasars, whose average and median $EW([\text{O III}])_{\text{rest}}$ values are 16.5 \AA and 13.0 \AA with the standard deviation of 14.6 \AA (the minimum and maximum values are 0.4 \AA and 73.0 \AA , respectively). Since the $EW_{\text{rest}}([\text{O III}])$ of SDSS J1707+6443 is 17.4 \AA (Table 2), the [O III] emission of SDSS J1707+6443 is not weaker than that of typical type 1 quasars. In Figure 7 we also show the frequency distribution of $EW_{\text{rest}}([\text{O III}])$ of SDSS quasars at $0 < z < 1$ (shown with filled circles with error bars in Figure 7). This is the distribution by Risaliti et al. (2011) who used the spectral measurements of SDSS DR5 quasars (Schneider et al. 2007) performed by Shen et al. (2008). The frequency distribution of $EW_{\text{rest}}([\text{O III}])$ of the $0 < z < 1$ SDSS sample (Risaliti et al. 2011; filled circles in Figure 7) is not significantly different from that of the higher- z ($1 < z < 4$) sample (solid histogram in Figure 7). The $EW_{\text{rest}}([\text{O III}])$ value of SDSS J1707+6443 is above the peak value of these distributions. Therefore we conclude that the $EW_{\text{rest}}([\text{O III}])$ of SDSS J1707+6443 is not significantly smaller than typical type 1 quasars in both low- z and high- z samples. Note that the actual relative [O III] $\lambda 5007$ strength of SDSS J1707+6443 with respect to the whole parent sample of type 1 quasars could be higher, because only [O III]-detected quasars (i.e., relatively strong [O III] emitters) are selectively shown in Figure 7.

The relatively large $EW([\text{O III}])_{\text{rest}}$ of SDSS J1707+6443 with respect to the global population of AGNs (Figure 7) sug-

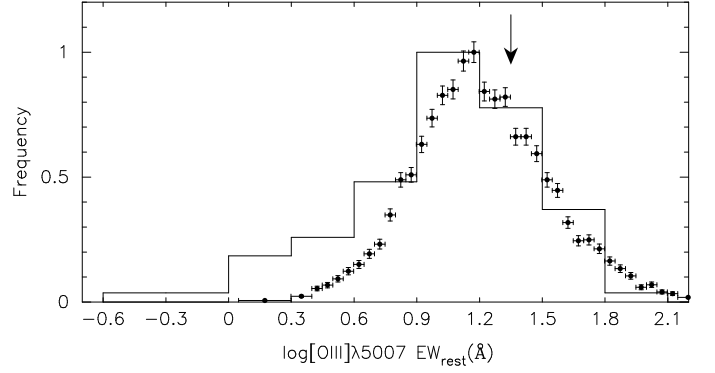


Fig. 7. Solid-line histogram denotes the compiled data of high- z type 1 quasars from the literature (McIntosh et al. 1999; Netzer et al. 2004; Marziani et al. 2009; Greene et al. 2010) quasars at $0 < z < 1$ (Risaliti et al. 2011). The frequency distributions of the two datasets are normalized by the number of objects in the bin at the peak of their frequency distributions. The arrow denotes the $EW_{\text{rest}}([\text{O III}])$ value of SDSS J1707+6443 measured in our work.

gests that the NLR in SDSS J1707+6443 is not characterized by a very high Z_{NLR} , when the photoionization model results shown in Figure 6 are taken into account. If a positive correlation between Z_{NLR} and Z_{BLR} is assumed, a lack of very high metallicity clouds in the NLR seems to be inconsistent with the strong broad nitrogen lines seen in the rest-frame UV spectrum of this object. Obviously these are two possible scenario to explain this inconsistency: (1) the strong UV broad nitrogen lines are due to a very high relative abundance of nitrogen in the BLR (with respect to non-N-loud quasars) rather than to an extremely higher BLR metallicity than non-N-loud quasars, or (2) the BLR metallicity is significantly higher than the NLR metallicity, which better represents the metallicity of the host galaxy. The relation between Z_{BLR} and Z_{NLR} is not yet well understood observationally, although Fu & Stockton (2007, 2008, 2009) reported that these two quantities should be related with each other, at least in low- z ($z < 0.5$) quasars. Whatever case (1) or (2) applies, the strong broad nitrogen lines of SDSS J1707+6443 are not consistent with the possibility that the host galaxy of this quasar is characterized by a very high metallicity. This implies that broad UV nitrogen lines of quasars are not (at least in some cases) a good tool to explore the chemical evolution of quasar host galaxies. Here we note that the spatial scale of NLRs is far larger than that of BLRs but not necessarily coincide with the spatial scale of their host galaxies; we show, however in the case of SDSS J1707+6443, the NLR spatial scale corresponds to the host-galaxy scale ($\sim \text{kpc}$ scale; see Section 4.3). Since quasar spectra are frequently used to investigate the chemical evolution at high redshifts, it will be crucial to examine whether the broad UV lines of quasars are good (or bad) tracers of the metallicity through more detail studies for larger samples of AGNs (see, e.g., Matsuoka et al. 2011).

4.2. Black Hole Mass and the Eddington Ratio

The results in the previous section suggest that the strong nitrogen emission in the BLR of SDSS J1707+6443 (or, possibly, N-loud quasars in general) is not indicative of high metallicity in the NLR and hence in the host galaxy. It is however important to verify whether the properties of the broad nitrogen lines

of SDSS J1707+6443 do follow the trends of the global population of high- z type 1 quasars. Recently Matsuoka et al. (2011) found that a high relative nitrogen abundance is seen in BLRs of high- $L_{\text{bol}}/L_{\text{Edd}}$ quasars. To examine whether SDSS J1707+6443 is consistent with the result of Matsuoka et al. (2011), we estimate M_{BH} and $L_{\text{bol}}/L_{\text{Edd}}$ based on the velocity widths of C iv in the SDSS spectrum and H β in our MOIRCS spectrum. Note that it is debatable whether C iv-based or H β -based estimation are more accurate (see, e.g., Netzer et al. 2004; Peterson et al. 2004; Denney et al. 2009). Therefore we use both emission lines to minimize possible systematic errors in the estimation of those parameters.

We derive M_{BH} for SDSS J1707+6443 by adopting the calibrations given by Vestergaard & Peterson (2006),

$$\log \frac{M_{\text{BH}}}{M_{\odot}} = 6.66 + 0.53 \log \left(\frac{\lambda L_{\lambda}(1350\text{\AA})}{10^{44} \text{ erg s}^{-1}} \right) + 2 \log \left(\frac{FWHM}{1000 \text{ km s}^{-1}} \right) \quad (1)$$

for C iv, and

$$\log \frac{M_{\text{BH}}}{M_{\odot}} = 6.91 + 0.50 \log \left(\frac{\lambda L_{\lambda}(5100\text{\AA})}{10^{44} \text{ erg s}^{-1}} \right) + 2 \log \left(\frac{FWHM}{1000 \text{ km s}^{-1}} \right) \quad (2)$$

for H β , where $FWHM_{\text{CIV}}$ and $FWHM_{\text{H}\beta}$ are the velocity width of the C iv and H β emission in full-width at half maximum, respectively. For estimating the Eddington ratio ($L_{\text{bol}}/L_{\text{Edd}}$), we adopt a bolometric correction of $L_{\text{bol}} = 9.26 \lambda L_{\lambda}5100$ (Shen et al. 2008). The derived values are $\log (M_{\text{BH}}/M_{\odot}) = 8.98$ and $\log (L_{\text{bol}}/L_{\text{Edd}}) = -0.53$ when using C iv, and $\log (M_{\text{BH}}/M_{\odot}) = 9.73$ and $\log (L_{\text{bol}}/L_{\text{Edd}}) = -0.21$ when using H β , respectively. Accordingly we adopt their means, $\log (M_{\text{BH}}/M_{\odot}) = 9.50$ and $\log (L_{\text{bol}}/L_{\text{Edd}}) = -0.34$. The derived Eddington ratio is relatively high (regardless of the adopted emission line for deriving M_{BH}), with respect to the frequency distribution of the Eddington ratio of SDSS type 1 quasars at similar redshifts (-0.48 ± 0.41 for 3144 quasars at $3.0 < z < 3.3$; Shen et al. 2011).

Our results is consistent with the finding of Matsuoka et al. (2011) that quasars with high accretion rate are not necessarily characterized by high metallicities, but are characterized by high nitrogen abundances (resulting in stronger nitrogen lines). In the scenario proposed by Matsuoka et al. (2011), the high black hole accretion is delayed by a few 100 Myr, relative to the main episode of star formation, when intermediate-mass stars have evolved and have enriched the ISM with nitrogen.

4.3. Emission-line Diagnostics in the Narrow Line Region

By combining the emission-line fluxes of [O iii] λ 5007 and [Ne iii] λ 3869 and the upper limit on the flux of [O ii] λ 3727, it is possible to investigate the properties of NLR gas clouds. This analysis enables us to study the physical properties of gas clouds in the host galaxy, because the NLR is extended on galactic scales in contrast to the BLR. Moreover, since the redshift of SDSS J1707+6443 ($z \sim 3$) corresponds to the peak of the global quasar activity (e.g., Richards et al. 2006), the gas properties of the quasar host galaxy are interesting to explore the interplay between AGNs and their host galaxies (i.e., the galaxy-SMBH coevolution). In this context, host galaxies of the N-loud quasars are particularly interesting since they may be in a special evolutionary stage, as mentioned in Section 4.2. Note that NLR emission lines other than [O iii] λ 5007 have only rarely been observed in luminous high- z type 1 quasars, thus not allowing a detailed investigation of NLR gas properties as done with our MOIRCS spectrum.

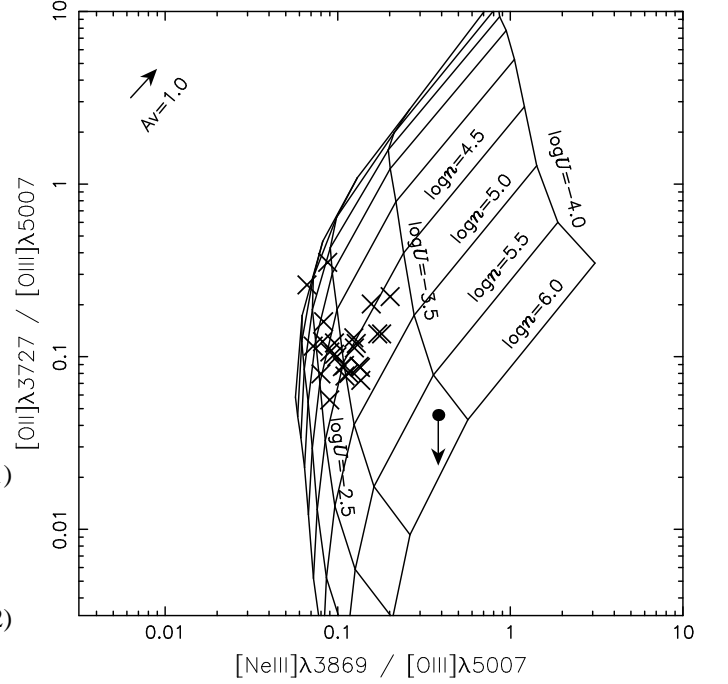


Fig. 8. Emission-line flux ratios of [O ii] λ 3727/[O iii] λ 5007 versus [Ne iii] λ 3869/[O iii] λ 5007. The filled circle with an arrow denotes the data of SDSS J1707+6443 (where the 3σ upper limit is adopted for the [O ii] flux), and cross symbols denote the data of the SDSS low- z ($z \sim 0.7$) type 1 quasars. The arrow at the upper-left corner in the panel is the reddening vector for observed data in the case of $A_V = 1.0$ mag, adopting the extinction curve of Cardelli et al. (1989). The grids are predicted emission-line flux ratios from Cloudy model runs, adopting $Z_{\text{NLR}} = 3Z_{\odot}$ and varying the hydrogen density and the ionization parameter.

In Figure 8, we place SDSS J1707+6443 on the diagram comparing the distributions of the emission-line flux ratios of [O ii] λ 3727/[O iii] λ 5007 and [Ne iii] λ 3869/[O iii] λ 5007. For comparison, we also plot the data of the SDSS DR7 type 1 quasar sample (Shen et al. 2011). We checked their SDSS archival spectra and selected quasars whose [O ii] λ 3727, [Ne iii] λ 3869, and [O iii] λ 5007 are significantly detected (i.e., $S/N > 10$). As a consequence we plot the emission-line flux ratio of 25 SDSS DR7 type 1 quasars, whose redshift range is $0.37 < z < 0.80$ with the median redshift of 0.65. Note that the average luminosity $\lambda L_{\lambda}(5100)$ for the 25 SDSS type 1 quasars is $10^{45.77} \text{ erg s}^{-1}$, that is ~ 1 dex lower than the corresponding value for SDSS J1707+6443 ($10^{46.66} \text{ erg s}^{-1}$). As shown in Figure 8, these 25 SDSS type 1 quasars show completely different flux ratios from SDSS J1707+6443; i.e., SDSS J1707+6443 shows a higher [Ne iii] λ 3869/[O iii] λ 5007 ratio and a lower [O ii] λ 3727/[O iii] λ 5007 ratio than SDSS type 1 quasars. These differences can be interpreted qualitatively if the NLR in SDSS J1707+6443 is characterized by a much higher average density than the NLRs in SDSS type 1 quasars at $z < 1$. This is because the [O ii] λ 3727 emission and partly [O iii] λ 5007 emission are suppressed by the collisional de-excitation process in high-density clouds with $n_{\text{H}} \sim 10^6 \text{ cm}^{-3}$ or higher, by taking the critical densities of these emission lines ($n_{\text{cr}}([\text{O ii}]) = (3.3\text{--}14) \times 10^3 \text{ cm}^{-3}$, $n_{\text{cr}}([\text{O iii}]) = 7.0 \times 10^5 \text{ cm}^{-3}$, and $n_{\text{cr}}([\text{Ne iii}]) = 9.7 \times 10^6 \text{ cm}^{-3}$) into account.

To investigate more quantitatively the differences in the NLR emission-line flux ratios between SDSS J1707+6443 and lower-

z type 1 quasars, we have performed Cloudy model runs for the parameter ranges of $n_{\text{H}} = 10^2 - 10^6 \text{ cm}^{-3}$ and $U = 10^{-4.0} - 10^{-1.5}$. Here we adopt a metallicity of $Z_{\text{NLR}} = 3 Z_{\odot}$, that is the typical value for NLR clouds (e.g., Groves et al. 2006), as we concluded in Section 4.1 that SDSS J1707+6443 is not characterized by very high Z_{NLR} . The other adopted parameters such as the input SED are the same as those described in Section 4.1. The results of the model runs are overlaid in Figure 8. The NLRs in the SDSS type 1 quasars are characterized by $n_{\text{H}} \sim 10^{4.5} \text{ cm}^{-3}$ and $U \sim 10^{-3.5} - 10^{-2.5}$, consistent with the parameter ranges inferred by previous studies (e.g., Nagao et al. 2001b). On the other hand, the NLR in SDSS J1707+6443 is characterized by a higher density, $n_{\text{H}} \sim 10^6 \text{ cm}^{-3}$ or more, although the ionization parameter is similar to that of type 1 SDSS quasars. These results are not sensitive to the adopted metallicity.

There are two possible scenarios to explain the distinct physical properties of SDSS J1707+6443. One possibility is considering a strong contribution from the inner wall of the dusty torus. Since the clouds located at the inner side of dusty tori are characterized by a high density and photoionized by the strong central continuum radiation, emission lines with high critical densities are radiated from these clouds (e.g., Murayama & Taniguchi 1998b; Nagao et al. 2000, 2001a). In the case that such high-density clouds at the inner dusty torus show strong NLR emission, its emission dominates the whole NLR emission and then low- n_{cr} emission lines become relatively weak, explaining the observed emission-line flux ratios of SDSS J1707+6443 (see also Murayama & Taniguchi 1998a; Nagao et al. 2001c). Another possibility is the existence of high-density clouds distributed on galactic scales (i.e., $\sim \text{kpc}$). Such a situation has been proposed for high- z and/or high-luminosity quasars (Netzer et al. 2004; see also Ho 2005), possibly related to violent star-formation activity in the host galaxy. Note that the latter scenario is not ruled out by the non-detection of the $[\text{O II}]\lambda 3727$ emission, which is used as an indicator of the star-formation rate (e.g., Gallagher et al. 1989; Kennicutt 1998; Ideue et al. 2009). This is because the $[\text{O II}]\lambda 3727$ emission is significantly suppressed in high-density H II regions when the density is higher than 10^4 cm^{-3} , even when the star-formation activity is very vigorous.

To distinguish between the above two scenarios, we estimated the typical distance of the ionized clouds from the nucleus in SDSS J1707+6443 by combining the estimates of n_{H} and U obtained in the previous section. Since we already estimated L_{bol} of SDSS J1707+6443 in Section 4.1, the number of the ionizing photon can be estimated for a given spectral energy distribution (SED). Here we assume the same SED used in the above photoionization model calculations (Mathews & Ferland 1987) and adopt the parameters of $n_{\text{H}} = 10^6 \text{ cm}^{-3}$ and $U = 10^{-3.5}$. Accordingly, we obtain the result that the typical distance of the clouds is $\sim 1.5 \text{ kpc}$, which is consistent with the latter scenario that considers high-density clouds in the host galaxy. Note that a smaller radius would have been derived by assuming a higher density, e.g., $\sim 15 \text{ pc}$ in the case of $n_{\text{H}} \sim 10^8 \text{ cm}^{-3}$. However this is not plausible, because $[\text{O III}]\lambda 5007$ emission would be significantly suppressed at such high-densities due to the collisional de-excitation. We thus conclude that there are high-density clouds at the kpc scale in the host galaxy of SDSS J1707+6443, that dominates the emission of the detected NLR lines. The torus scenario is disfavored also in terms of the observed NLR velocity width. By assuming that the SMBH is the dominant source of the gravitational potential field in the nucleus of SDSS J1707+6443, the expected velocity width of $[\text{O III}]\lambda 5007$ is $\sim 1300 \text{ km s}^{-1}$ if the most of this emission arise

at $\sim 15 \text{ pc}$ from the SMBH with $M_{\text{BH}} = 10^{9.78} M_{\odot}$ (see Section 4.2). This expected velocity width is apparently larger than the observed width (617 km s^{-1} ; Table 2), suggesting that most of the $[\text{O III}]\lambda 5007$ emission in this object arises at larger spatial scales that probably corresponds to the scale of its host galaxy.

Here we briefly discuss the implication of the extended ($\sim \text{kpc}$) dense ($\sim 10^6 \text{ cm}^{-3}$) gas clouds in SDSS J1707+6443 inferred from emission-line diagnostics. It is well known that such dense clouds in galaxies are closely related to the star-forming activity. Plume et al. (1997) reported their observations of the carbon monosulphide (CS) molecule for >100 Galactic H II regions and showed that their typical density is $n \sim 10^6 \text{ cm}^{-3}$ (but less than $2 \times 10^7 \text{ cm}^{-3}$). Such a high density for clouds in H II regions has been inferred also from some other molecular-line radio observations (e.g., Hofner et al. 1996; Bergin et al. 1996; Lada et al. 1997). Although such a high density is not the "typical" density for H II region clouds, the star-formation efficiency of the denser clouds is actually much higher than less dense clouds. This is suggested by, e.g., a clear positive correlation between the $L_{\text{HCN}}/L_{\text{CO}}$ ratio (i.e., the dense-gas fraction) and the $L_{\text{IR}}/L_{\text{CO}}$ ratio (i.e., the star-formation efficiency) that is seen in H II region clouds (e.g., Gao & Solomon 2004). Therefore high-density clouds generally dominate the star formation, when they exist (see also, e.g., Wang et al. 2011). A similar situation is also seen in high- z galaxies (e.g., Gao et al. 2007; Riechers et al. 2007, 2010, 2011) but with a significant difference compared to low- z galaxies; that is the spatial extension of vigorous star-forming regions. The spatial scale of star-forming regions in high- z actively star-forming galaxies (such as sub-millimeter galaxies) extends up to $\sim \text{kpc}$ scales or more (e.g., Chapman et al. 2004; Walter et al. 2009), that is different from low- z actively star-forming galaxies such as ultra-luminous infrared galaxies (e.g., Downes & Solomon 1998; see also, e.g., Iono et al. 2009). These pictures well match with our results in the sense that dense gas clouds ($\sim 10^6 \text{ cm}^{-3}$) are distributed at the $\sim \text{kpc}$ scale in a galaxy at $z \sim 3.2$. Therefore, given these considerations, we speculate the existence of vigorous star-forming activity in SDSS J1707+6443.

5. Summary

To assess the physical and chemical properties of the host galaxies of N-loud quasars, we analyzed the MOIRC near-infrared spectrum of SDSS J1707+6443, at $z=3.16$, obtaining the following results:

- This N-loud quasar shows strong $[\text{O III}]\lambda 5007$ emission. Since photoionization models predict weak $[\text{O III}]\lambda 5007$ emission when the NLR metallicity is very high ($> 5 Z_{\odot}$), the detected strong $[\text{O III}]\lambda 5007$ emission suggests that the NLR in this object is not characterized by very high metallicities.
- The UV nitrogen lines from BLRs are not ideal tools to discuss the chemical evolution of quasar host galaxies, because whatever the origin of strong broad nitrogen emission (high Z_{BLR} , or high relative abundance of nitrogen with an ordinary metallicity), Z_{BLR} are likely unrelated with the metallicities of the host galaxies.
- The Eddington ratio of SDSS J1707+6443 is moderately high ($L_{\text{bol}}/L_{\text{Edd}} = 10^{-0.26}$), as derived from single-epoch M_{BH} estimates using both C IV and H β emission lines. This is consistent with the discovery by Matsuoka et al. (2011) that a high relative nitrogen abundance is associated to quasars with high Eddington ratios.

- The flux ratio of $[\text{O II}]\lambda 3727/[\text{O III}]\lambda 5007$ in SDSS J1707+6443 is significantly lower than that in lower- z type 1 quasars, indicating that NLR clouds in SDSS J1707+6443 are characterized by much higher densities than those in other type 1 quasars.
- Photoionization models suggest that those high-density clouds (with $n_{\text{H}} \sim 10^6 \text{ cm}^{-3}$) are located at the kpc scales in the host galaxy of SDSS J1707+6443, and we speculate that this might be related to the vigorous star-formation activity in the host galaxy.

These results possibly reveal a strong connection between the high AGN activity (characterized by a high Eddington ratio) and the vigorous star-formation activity suggested by the kpc-scale distribution of dense gas cloud in a N-loud quasar, SDSS J1707+6443. Currently it is not clear whether such a connection is seen also in other N-loud quasars generally or not, given a paucity of detailed near-infrared spectroscopic studies for N-loud quasars. Therefore it is highly interesting to examine whether such a situation is common in other N-loud quasars through further near-infrared spectroscopic observations for a large sample of N-loud quasars.

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